

Advances in Solar Physics

The Solar Physics Section of EPS held its triennial meeting *Advances in Solar Physics* this year in Catania, from 11-15 May. The meeting was attended by 140 participants, including over 20 from eastern Europe and several from the United States. The meeting comprised seven scientific sessions and was opened by an outstanding talk titled 'The Sun Today' by J.-C. Pecker of the Collège de France, Paris.

The meeting aimed to bring together solar physicists studying on one hand the solar interior, and on the other hand the solar atmosphere. In order to link our understanding of fundamental processes such as convection and magnetic field generation to that of energy release and transport in the sun's outer layers (i.e., radiation, mass motion, heating, and particle acceleration), it is of crucial importance that these two communities combine their efforts. Comparisons of reconnection and acceleration processes between laboratory and natural plasmas appear to be a promising approach for the future.

The meeting's first session dealt with the problem of the lack of neutrinos in the frame of the new instruments in Italy (notably those at the Gran Sasso laboratory) and in the former eastern block. The conversion of neutrinos into particles with different mass may be a possible solution.

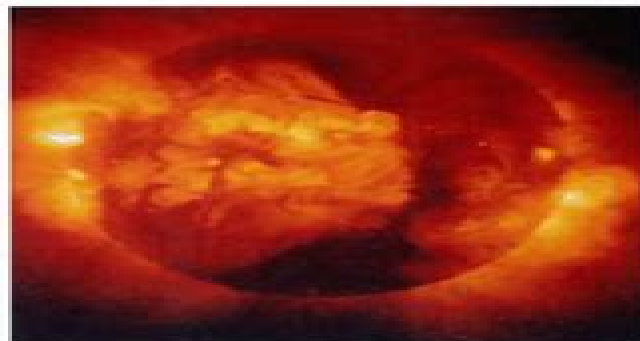
The different zones inside the sun have been addressed specifically with respect to their rotation rates, and this may be important for explaining the duration of the solar cycle. The dynamo effect certainly operates at the base of the convective zone (CZ) in a sheared region located between the solar core with its rigid rotation and the overlying CZ,

with its differential rotation. Recent imaging observations of the corona in the radio domain with the Nancyair radio heliograph, and in the soft-X-ray domain with the Japanese satellite YOHKOH, have emphasized the dynamic behavior of the large, evolving loops. Their constant evolution is presumably due to convective motions at foot-prints. The heating is connected with the magnetic field: Alfvén waves or reconnection of magnetic field lines.

Two sessions were devoted to the flare physics. These phenomena represent a large release of energy — up to 10^{32} ergs. Using high temporal and spatial resolutions we can detect a succession of bursts and kernels which seem to indicate that some fragmentation of the energy release is occurring over very short time scales. This was shown using the radio-heliograph of Nancyair and the new gamma-ray detector aboard the Russian satellite GRANAT.

The acceleration of particles is still a controversial topic. Are the particles accelerated slowly before the flare, trapped in the corona, and then suddenly dumped into the photosphere; or are they all accelerated impulsively at the impulsive phase of the flare; or are some accelerated by post-flare shocks as they move through the corona?

The last session was concerned with instrumentation: ground-based instruments (in the Canary Islands and at Pic du Midi) and space instruments (the future ESA-NASA



An image of the sun taken in the soft X-ray region from the YOHKOH spacecraft. It shows highly ionized plasma emitting at wavelengths between 2 and 20 Å. Small flares can be seen on the east and west limbs (right and left, respectively), and an extensive loop structure associated with an active region is seen just above the centre. The coronal hole that stretches from the north to south poles persisted for several months.

satellite SOHO). During the last decade, improvements have been made in spatial resolution for intensity and velocity measurements. The French-Italian telescope THEMIS will provide in 1996 magnetic field vector measurements with a spatial resolution four times better than those now available.

The database generated from new ground-based and space instruments available to the European solar community was discussed

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